

Spectroscopy of bright quasars: emission lines and internal extinction[★]

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ABSTRACT

Aims. The main purpose of this work is to improve the existing knowledge about the most powerful engines in the Universe – quasars. Although a lot is already known, we still have only a vague idea how these engines work exactly, why they behave as they do, and what the relation is between their evolution and the evolution of their harboring galaxy.

Methods. Methods we used are based on optical spectroscopy of visually bright quasars, many of which have recently been discovered as X-ray sources, but eventually missed in color-selected surveys. The spectra typically cover the 4200–7000 Å region, allowing measurements of the characteristics of the hydrogen lines, the FeII contribution, and other lines of interest.

Results. We present accurate redshift estimates and Seyfert type classification of the objects. We also show that the contribution of the host galaxy to the optical continuum is non-negligible in many cases, as is the intrinsic AGN absorption. Consequences of not correcting for those factors when estimating different quasar parameters are discussed. We also find some evidence of a non-unity slope in the relation between the internal extinction based on the Balmer decrement and the one on the optical continuum slope, implying, if further confirmed, the intriguing possibility that some absorbing material might actually be located *between* the continuum source and the broad-line region.

Key words. galaxies: Seyfert – galaxies: quasars: general – galaxies: quasars: emission lines

1. Introduction

The generally accepted paradigm about the nature of quasars (or active galactic nuclei; AGN) invokes a supermassive black hole, surrounded by an accretion disk, where most of the emitted energy is generated. In the vicinities of this disk, there must be a region, producing the broad emission lines (BLR), traditionally presented as a system of clouds, moving in the gravitational field of the black hole. Farther out from the center should be a region producing narrow lines (NLR), which may be the only visible optical lines in the spectra, if the central region is hidden from the observer behind a thick torus, as the unification models explain the AGN type I/II differences (Antonucci 1993, for a review). Studying the narrow lines is important, since they are fairly sensitive to the shape of the ionizing continuum, and especially to the unobserved EUV range, which, if restored properly, may help to test the general paradigm or to model the type of accretion disk (e.g. standard thin disk, ADAF, etc.). One of the least clear points in this picture so far appears to be the exact geometry (e.g. flat, spherical, conical), structure (system of clouds, irradiated accretion disk, etc.), and kinematics (inflow, outflow, Keplerian rotation) of the BLR matter. The BLR should play an important role in feeding the quasar; therefore, knowledge of

its nature may reveal why some galactic centers are active and others are not, in case a supermassive black hole is likely to be present in both types (Magorrian et al. 1998).

The best way to probe the broad-line region might be through its optical emission-line spectrum. Although a single spectrum is certainly not sufficient to answer all the questions, studying the line profiles, ratios, and other properties in many objects seems to be one way of shedding some light on the problem. A promising new approach to systematizing quasar properties and differentiating between quasars is based on principle component analysis (PCA), where the quasar diversity can be expressed as a linear combination of a few eigenvectors of the correlation matrix, constructed from different measured quantities of a sample of quasars. Many of these measures are based on the optical spectra (Boroson & Green 1992; Sulentic et al. 2000), such as the width of the broad lines, equivalent width ratios. It is proposed that the main eigenvector (EV1), i.e. the main driver of the quasar diversity, is the accretion rate ratio (Marziani et al. 2001), one of the defining characteristics of the AGN. On the other hand, another important characteristic of the central engine, the black hole mass, is often calibrated through empirical relations, like the “width of the broad lines, $FWHM$ – the monochromatic continuum luminosity, $\lambda L_{\lambda 5100 \text{ Å}}$ ”, (Kaspi et al. 2005), “the black hole mass – bulge luminosity or bulge velocity dispersion” (Magorrian et al. 1998; Ferrarese & Merritt 2000), etc., since the only direct method known, the reverberation mapping (see Peterson & Horne 2004, for a review), is observationally very expensive. It is clear that a large sample of quasars with relatively good spectra is needed for better understanding

[★] Based on observations obtained with the 2-m telescope of the Rozhen National Astronomical Observatory, which is operated by the Institute of Astronomy, Bulgarian Academy of Sciences, and with the 1.3-m telescope of the Skinakas Observatory, Crete, Greece; Skinakas Observatory is a collaborative project of the University of Crete, the Foundation for Research and Technology – Hellas, and the Max-Planck-Institut für Extraterrestrische Physik.

Table 1. Log of observations.

Object	RA J2000	Dec J2000	z	$\sigma(z)$	A_V	Date	Instr.	Exp.	S/N
2MASXJ005050+3536	00 50 51	+35 36 43	0.0585	0.0004	0.14	08.09.2004	Rz 2.0	2700	36
MCG+08.17.060	09 13 46	+47 42 00	0.0524	0.0005	0.05	19.03.2004	Rz 2.0	2×2400	10
RXS J11401+4115	11 40 03	+41 15 05	0.0717	0.0002	0.06	22.06.2005	Sk 1.3	2×1200	12
PG 1211+143	12 14 18	+14 03 13	0.0815	0.0003	0.11	23.06.2005	Sk 1.3	2×1200	27
RXS J12308+0115	12 30 50	+01 15 21	0.1183	0.0005	0.06	24.06.2005	Sk 1.3	1200	25
RXS J16312+0955	16 31 16	+09 55 58	0.0917	0.0005	0.18	17.08.2004	Sk 1.3	2×2700	35
RXS J17233+3630	17 23 25	+36 30 25	0.0400	0.0003	0.16	29.06.2005	Rz 2.0	1800	25
IRAS 18423+2201	18 44 30	+22 04 28	0.0464	0.0004	0.69	11.08.2005	Sk 1.3	2×2400	15
NPM1G+27.0587	18 53 04	+27 50 28	0.0620	0.0004	0.49	10.08.2005	Sk 1.3	2×1800	31
						04.05.2005	Rz 2.0	1800	29
RXS J20440+2833	20 44 04	+28 33 09	0.0492	0.0005	1.03	28.08.2005	Sk 1.3	2400	38
						30.06.2005	Rz 2.0	1800	33
NPM1G–05.0589	21 03 38	–04 55 40	0.0575	0.0002	0.27	18.08.2004	Sk 1.3	2×2700	31
RXS J21240–0021	21 24 02	–00 21 58	0.0617	0.0003	0.15	17.08.2004	Sk 1.3	2×2700	36
NPM1G+24.0470	21 39 41	+24 24 18	0.0390	0.0005	0.23	10.08.2005	Sk 1.3	2400	26
RXS J21592+0952	21 59 12	+09 52 42	0.1003	0.0002	0.21	19.08.2004	Sk 1.3	2×2700	39
RXS J22027–1304	22 02 45	–13 04 53	0.0391	0.0003	0.14	28.08.2005	Sk 1.3	2×2400	22
RXS J22160+1107	22 16 04	+11 07 26	0.0514	0.0005	0.21	10.08.2005	Sk 1.3	2400	16
RXS J22287+3335	22 28 46	+33 35 08	0.0906	0.0004	0.29	11.08.2005	Sk 1.3	2×2400	35
NPM1G–04.0637	22 53 11	–04 08 49	0.0257	0.0002	0.14	11.08.2005	Sk 1.3	2×2400	24

the importance of EV1 correlations for the quasar physics and to better calibrate different quasar empirical relations.

Nowadays we see that many bright broad-line quasars have been missed in color-selected samples (e.g. Palomar-Green et al. 1983), mainly because of their reddened colors, due to significant host galaxy contribution and/or intrinsic absorption. Being mostly radio-quiet, many of those objects have only recently been discovered and identified as quasars through X-ray surveys. Studying such objects in detail is especially important for preventing introduction of different biases into our knowledge about the overall AGN population, its properties, and evolution. Furthermore, these partially obscured objects (reddened S1, S1.8, S1.9 types, etc.) may provide an important link between the “pure” type I AGN and the obscured population (type II/Seyfert 2), especially taking into account the building evidence that the latter may comprise a significant part of the entire AGN population (Zakamska et al. 2003). This paper should be considered as a part of a larger continuing effort to obtain spectra of relatively good quality for most of the nearby, visually bright quasars, including many discovered in recent years. In a subsequent paper, we will concentrate on the connection between the quasar host galaxies and different AGN characteristics.

The paper is organized as follows. In the next two chapters we describe the observations and present the main results, including individual notes on the objects. Discussion is presented in Sect. 4 and the summary, in Sect. 5.

2. Observations and reductions

2.1. Objects

We did not apply any formal criteria to select the objects. Our intention was mostly to observe type I objects from the Véron-Cetty & Véron catalog (2003, hereafter VCV03) with no previously published spectra of good quality and resolution. The objects were bright enough to be accessible with 1.5–2-m class telescopes (Sect. 2.2), i.e. – 14–15th magnitude. Most of these objects were discovered initially as X-ray sources (RASS catalog) and identified as quasars/Seyfert galaxies only recently. One can naturally expect that these objects will appear to be

predominantly radio quiet, as in our case. A few quasars, like PG 1211+143, are well-studied objects; however, they have been observed to be able to better assess our measurement errors and the quality of spectral calibration and to look for possible variations. The most common names of the objects and their J2000 coordinates, as given by VCV03 are shown in Table 1 (Cols. 1, 2). The next columns in the table contain our measured redshifts with the errors (Sect. 2.3 for details) and the Galactic absorption.

2.2. Observational data

The last four columns in Table 1 show the observation date, the instrument (Rz 2.0 is the Rozhen telescope, Sk 1.3 is the Skinakas telescope, see the text below), the exposure time in seconds and the signal-to-noise ratio. The spectra of the objects were obtained with two instruments – the 2-m RCC telescope of the Rozhen National Observatory, Bulgaria, and the 1.3-m RC telescope of Skinakas Observatory, Crete. The Rozhen telescope is equipped with a Photometrics AT200 CCD camera (1024×1024 pixels) and the focal reducer FoReRo attached to the RC focus of the telescope. A grating prism (300 lines/mm) and $130 \mu\text{m}$ slit were used. The Skinakas telescope is equipped with a 2000×800 ISA 612 CCD attached to a focal reducer. A grating prism of 1300 lines/mm and a slit of $160 \mu\text{m}$ were used there.

A standard HEAR lamp or night sky lines were used for the wavelength calibration, allowing a wavelength accuracy typically better than 2–3 Å (Table 1). Suitably located spectroscopic standards were used for the flux calibration. The typical signal-to-noise ratio of the spectra is 25–30 and the spectral resolution – 5–10 Å (Table 1).

2.3. Reductions

We used standard IRAF routines to perform wavelength and flux calibrations and to extract 1D spectra of the objects. After that, each spectrum has been corrected for the atmospheric A and B-band absorption, which in a few cases, unfortunately, appeared very close to strong emission lines of interest. The results

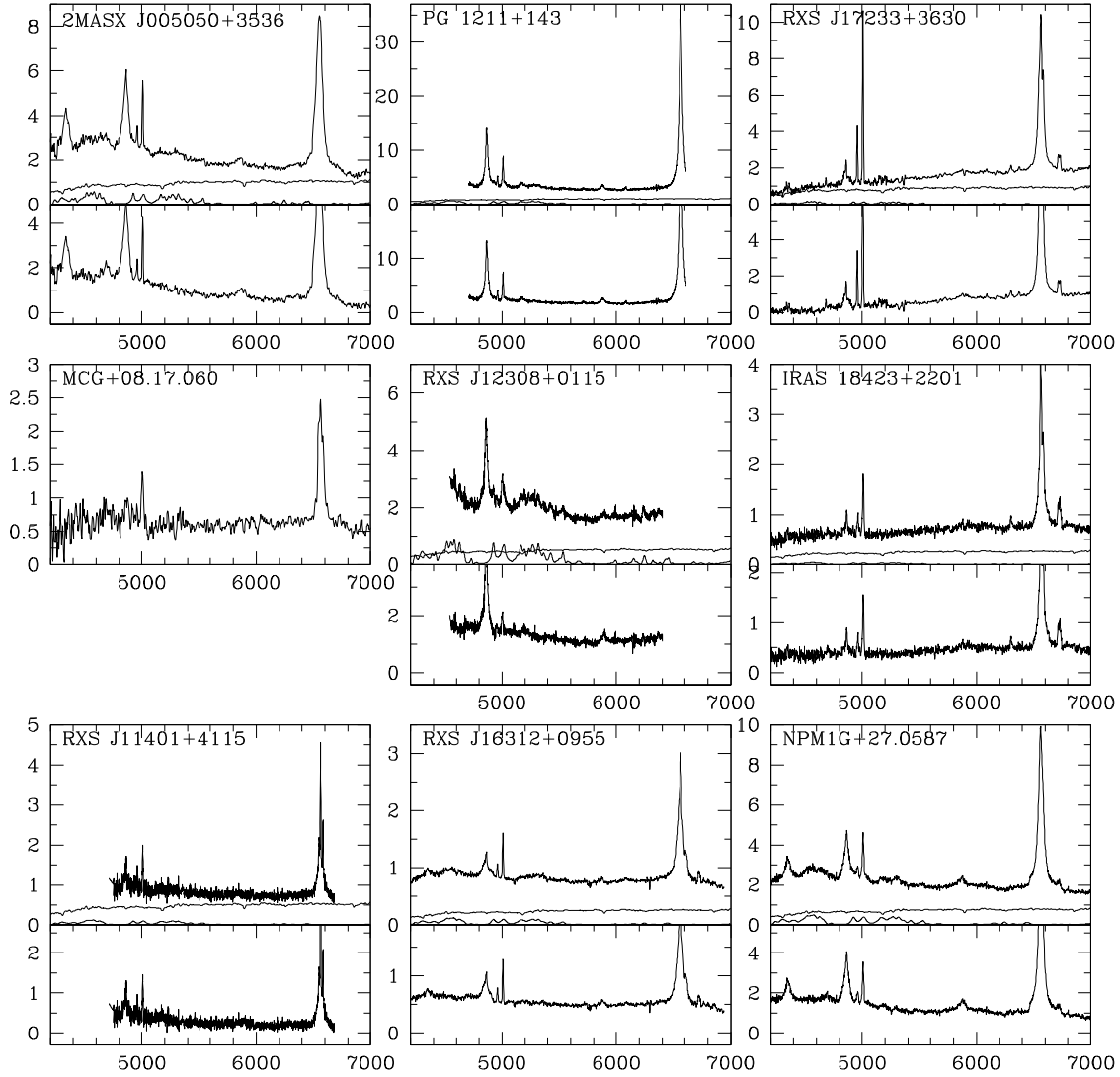


Fig. 1. Optical spectra of all objects from our sample. All spectra are corrected for atmospheric absorption bands, Galactic absorption and are deredshifted. The upper part of each box displays the original spectrum (thick line) and the FeII and starlight templates to be subtracted from it (thin lines). The lower part shows the resulting intrinsic AGN spectrum. All fluxes are in units of $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$.

from this correction are rather satisfactory for the majority of the objects. All spectra were also corrected for the Galactic absorption (Schlegel et al. 1998), which although typically rather small ($A_V \approx 0.1\text{--}0.2$ mag), turned out to be significant for some objects – in order of a magnitude (Table 1). The [OIII] $\lambda 5007$ line was mostly used to calculate the redshift and to perform the redshift correction, since it was often the most prominent narrow line in the spectrum. In a few spectra, where [OIII] was weak (and likely to be blue-shifted, e.g. Zamanov et al. 2002), we used the narrow top of H β and/or H α . The redshifts are listed in Table 1. We also show typical errors, which are based on the quality of the spectrum and on how well redshifts of different lines agree in between. Our typical redshift error is ≈ 0.0005 or less, and the standard deviation between our redshifts and those from VCV03 is ≈ 0.001 .

3. Results

3.1. Host galaxy and FeII contributions

The redshift-corrected spectra of the objects included in this study are shown in Figs. 1 and 2. To be able to measure the

line profiles and the intrinsic power-law continuum accurately, two more corrections had to be made. The first one is the subtraction of the host galaxy (starlight) contribution to the continuum. The host galaxy signatures are clearly seen in many objects that finally turned out to be S1.8/S1.9 type galaxies and, to some extent, in S1 types as well. One can approximately estimate the host galaxy contribution by the strength of the stellar absorption lines, such as the CN g-band (around 4300 Å), MgB triplet (around 5172 Å), Ca+Fe (around 5269 Å), and NaD lines (5892 Å). We subtracted a scaled S0 template from each spectrum, made publicly available by Kinney et al. (1996); see Hao et al. (2005) and Strateva et al. (2005) for an alternative approach. The S0 galaxy was arbitrary chosen at first, because we have little knowledge about the exact host galaxy type in most cases, and second, because many spectra are rather noisy to be able to precisely distinguish the templates, and finally, because the differences between the templates are actually not that large. The scaling of the template was chosen in such a way that no signs of absorption lines are seen in the resulting spectra. To judge the strength of the host galaxy contribution, the characteristic shape of spectra in the 4200–4600 Å range, where the

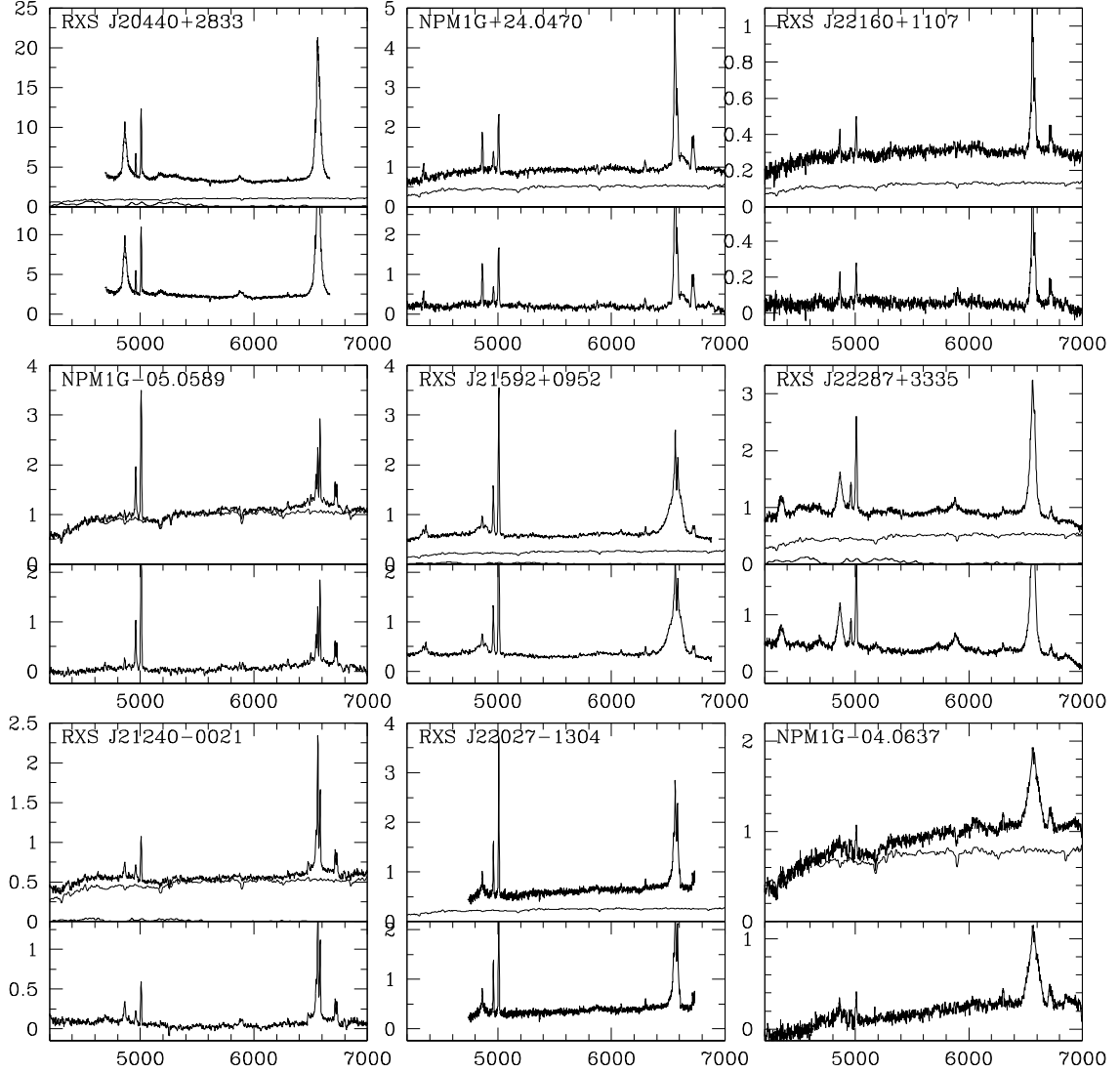


Fig. 2. See Fig. 1.

host galaxy continuum changes significantly, is also helpful. Of course, such an approach is prone to subjectivity, which in addition to possible systematic errors (mismatched galaxy type, possible extra starburst component, possible problems with the flux calibration of the spectra, noise, etc.), does not allow us to make any strong statements about the exact value of the host galaxy contribution. It can, however, give some good idea about how much this contribution can be, and indeed, our data suggest that it is significant in many cases (Figs. 1, 2; Sects. 3, 4). For a few objects, the host galaxy subtraction helped to reveal broad wings of the $H\alpha$ and/or $H\beta$ lines, clearly seen only on the remaining flat continuum after the subtraction, which is very important for the correct classification of the objects.

Another “contaminating” contribution is from the FeII lines. Of course, being intrinsic to the BLR, these lines cannot be considered as a contamination, but their subtraction is necessary for two reasons. First, it allows better measurements of lines like $\text{HeII}\lambda 4686$, $[\text{OIII}]\lambda 4959, 5007$, and to some extent, $H\beta$. After the FeII template subtraction, a broad feature, possibly a blend of $[\text{FeVII}]\lambda 5160$, $[\text{FeVI}]\lambda 5177$, and $[\text{NI}]\lambda 5200$ is clearly seen in many objects at a level roughly about 20% from the FeII flux (Table 3). This feature is otherwise completely hidden in the template. It may, indeed, be partially due to a poor host galaxy

subtraction. Second, the EV1 correlations reveal the significance of the FeII emission, where $R_{\text{Fe}} = F(\text{FeII})/F(H\beta_{\text{broad}})$ is one of the most important parameters used there. Therefore, the FeII contribution has to be estimated as precisely as possible. We used a broadened and scaled FeII template of I Zw 1 as is usually done (see for instance Marziani et al. 2003, for more details on the FeII subtraction). Again, no formal criterion was used to estimate the iron strength, and it has to be considered as a best guess with an error typically of about 20% or higher for the noisier spectra. One can judge how successful the FeII subtraction was by the flatness of the continuum around $H\beta$ (Figs. 1, 2, the lower panel of each box). The FeII emission flux ratio (R_{Fe}) is given in Table 2. The FeII flux is obtained by integrating the template between 4434 and 4684 Å, as in Boroson & Green (1992). This is only ~25% of the total FeII emission between 4000 and 7000 Å, so for many objects, iron has to be considered as producing the most prominent line features in optics, not hydrogen. The order of the host galaxy and the FeII subtractions was not the same for all the objects; we started with the template, whose contribution seemed more significant. Figures 1 and 2 show all the spectra. The upper panel of each box displays the original spectrum and the templates to be subtracted from it. The lower panel shows the spectrum after the subtractions, used further for

Table 2. Broad-line measurements.

Object	Flux H β	<i>FWHM</i> H β	Fit H β	R_{Fe}	Flux H α	<i>FWHM</i> H α	Flux HeII λ 686	Flux HeI λ 5876	$\log \lambda L_{\lambda}$ 5100 Å	a_v	$\log M_{\text{BH}}$
2MASXJ005050+35	180	3690	G	0.49	460	3060	39	37	44.0	2.7	7.5
MCG+08.17.060	?	?		?	92	2740	?	?	43.6	?	?
RXS J11401+4115	21	3130	G?	1.05	72	2290			43.7	1.4	7.2
PG 1211+143	415	1810	L	0.30	1400	1280		38	44.6	-0.5	7.3
RXS J12308+0115	106	2120	L?	1.25			2	?	44.7	-0.6	7.5
RXS J16312+0955	23	4380	G	0.96	113	2290	3	5	44.1	-1.4	7.7
RXS J17233+3630	55	4920	G	≤ 0.5	370	2600	15	21	43.2	-5.8	7.2
IRAS 18423+2201	14	3810	G	≤ 0.5	90	?	4	10?	43.3	-3.4	7.1
NPM1G+27.0587	113	3870	G	0.78	460	2740	25	75	44.1	0.4	7.6
RXS J20440+2833	268	2500	L?	0.46	837	1690		44	44.2	-1.0	7.3
NPM1G-05.0589				-	<30	11450?	2		-	-	-
RXS J21240-0021	6	4320	G	<0.1	<24	6850?	5	3?	42.8	-0.2	6.8
NPM1G+24.0470				-	27	6220	0	3	42.9	-0.7	-
RXS J21592+0952	20	6950	G	0.36	133	7400	2	?	43.9	-2.0	8.0
RXS J22027-1304	15	4750	G	<0.1	50	2560		7	43.1	-3.2	7.1
RXS J22160+1107	3	3310	G	-	13	3650?	4	?	42.6	-0.4	6.4
RXS J22287+3335	36	3630	G	0.61	130	2610	12	25	43.9	-0.7	7.4
NPM1G-04.0637	14	7250	G	<0.1	84	4750	9?	0	42.2	-7.6	6.9

Table 3. Narrow-line flux measurements in units of $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ Å}^{-1}$.

Objects	H β 4861	[OIII] 5007	[FeVI] 5177	[FeVII] 6087	[OI] 6300	H α 6563	[NII] 6548	[NII] 6583	[SII] 6716+31	AGN type
2MASXJ005050+3536	11	30			11				0	S1
MCG+08.17.060	?	17			3?	4	3?	5	4	S1
RXS J11401+4115	4.8	7.8	5		2	15	2.2	8.5		(LL)S1
PG 1211+143	25	45	23	74	0					NLS1
RXS J12308+0115	8?	9.6								(NL)S1
RXS J16312+0955	1.6	8			?				4	S1
RXS J17233+3630	11	94	3.1	3.5	8	32	?	18	20	S1
IRAS 18423+2201	4.3	15			5.2	21?			13	S1 (L1)
NPM1G+27.0587	8.2	27	5		5				13	S1
RXS J20440+2833	10	70	30		9					S1
NPM1G-05.0589	2.6	30			2	10	7	16	9.5	S1.9?
RXS J21240-0021	2.5	6			1?	10?	4.6	9.1	6.5	S1 (L1S1, L1)
NPM1G+24.0470	12	19			3.5	60	10?	29	19.8	S1.9? (L1, HII)
RXS J21592+0952	2.8	33		2	2.6	11	3	7	4.6	S1
RXS J22027-1304	2.2	27		2	2.3	9	5?	9	5	S1.8
RXS J22160+1107	1.1	2.4				6?	2.7	5.8	3.6	S1.8 (L1S1)
RXS J22287+3335	2.6	22	5	3	2.8				4.2	S1
NPM1G-04.0637	2?	4			3				7	S1

the emission line and continuum measurements (Sect. 3.2). We do not try to fit a power law continuum to the resulting quasar spectrum (e.g. Dietrich et al. 2005), since there is no guarantee that a single power-law describes the entire AGN optical continuum (although we do measure a local slope, Sect. 3.5).

3.2. Emission-line measurements

Fluxes (in units of $10^{-15} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ Å}^{-1}$) of various lines as measured from the spectra are shown in Tables 2 and 3. All widths are in km s^{-1} , λL_{λ} is in erg s^{-1} , and the black hole masses (Table 2, the last column, see also below) are in Solar units. In Tables 2 and 3 question marks indicate insecure line measurements or lines, too noisy to measure accurately. When no line is seen in the spectrum, the corresponding box is empty. The typical errors of the optical slopes and the central masses (Table 2, the last two columns) are estimated to be ± 0.5 .

The narrow lines (Table 3) are measured after fitting with a Gaussian. [FeVI] λ 5177 is probably blended with [FeVII] λ 5160 and [NI] λ 5200. The typical errors of the narrow-line

measurements are $\sim 30\%$. No narrow line widths are shown, since the widths are often determined by the instrumental width (or we cannot verify the opposite). We also do not show the equivalent widths (EW), due to the uncertainty of the continuum level, often dominated by the host galaxy (see above), in which case EW hardly has any meaning. Where present, the broad lines (H α and H β) were measured after an attempt to subtract the narrow component on top. We succeeded in isolating a narrow component in practically all H β profiles, where in some cases inflection was apparent, while in other – we subtracted our “best guess” for the narrow component, which was supposed to have the same width as the [OIII] lines. The errors in the latter case can be, of course, significant (as much as 50% – much more than the typical error of [OIII], which is about 10%), at least because a narrow top does not necessarily mean an NLR emission. A similar approach was used for H α lines; however there, the blended [NII] λ 6548, 6583 Å lines could not be separated from the H α profile in many cases, especially when the H α profile itself was broad enough (see Sect. 3.5 for details on the individual spectra). In those cases, only a broad

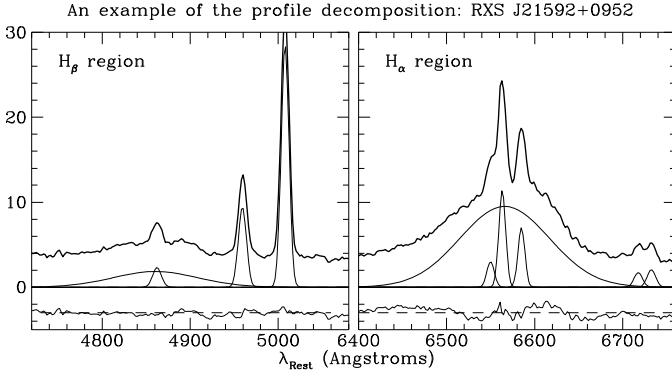


Fig. 3. An example of profile decompositions for one of the objects (RXS J21592+0952). The thick line (*upper panels*) represents the original spectrum (after galaxy and FeII subtractions). The thin lines are the fitted lines, where only Gaussians are used. The residuals are shown in the bottom panels. *The left panels* are for the H β region ([OIII] and H β broad and narrow components are shown) and *the right panels* – the H α region ([SII], [NII] and both components of H α are shown). One sees that a single function (Gaussian or Lorentzian) does not always model the broad emission perfectly (which is not the purpose of this paper), as seen from the residuals for the H α region. The fluxes are in $10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$.

component was measured, and one should be aware about some possible overestimation of both – the flux and the $FWHM$ of H α (actually $FWHM$ can also be underestimated if the narrow component of H α is much stronger than the [NII] lines). Where possible, the narrow component of H α and NII lines were deblended and they, together with the broad component, were measured separately. The parameters of the broad components ($FWHM$, flux) were measured after fitting with Gaussian or Lorentzian functions. The choice was made upon the goodness of the fit and is indicated in Table 2. For most of the objects, one of these two possibilities worked quite well (Sect. 4.1). The measurement errors of the $FWHM$ and of the flux of the broad components are typically 10–20%, estimated after different fits and continuum levels were attempted.

Figure 3 shows an example of profile decompositions for one of the objects (RXS J21592+0952). Our sample seems to confirm that the narrower-line objects (and eventually those with stronger FeII emission) tend to be better fitted by a Lorentzian, and vice versa (Sulentic et al. 2000; Véron-Cetty et al. 2001). See however the Sect. 4.1 for comments on this issue. Table 2 also contains the monochromatic flux at 5100 Å, often used for BLR size (R_{BLR}) and central mass (M_{BH}) estimates. The value of $\lambda L_{\lambda 5100}$ is measured on the resulting spectrum after the FeII and starlight corrections and should be considered as a measure of the intrinsic quasar optical continuum. Due to the uncertainties of those corrections, the uncertainties of $\lambda L_{\lambda 5100}$ have to be in the 30–40% range. Knowing the width of the lines and $\lambda L_{\lambda 5100}$, we can estimate the black hole mass for each S1 object (Table 2), assuming

$\log(M_{BH}) = 0.69 \log(\lambda L_{\lambda}) + 2 \log(FWHM(H\beta)) - 30.0$, obtained following Kaspi et al. (2005). No corrections for the intrinsic absorption have been made, and if significant, the absorption can lead to underestimating the continuum level, and to underestimating of M_{BH} up to a magnitude or even more (Sect. 4). An alternative calibration (Bentz et al. 2006), taking the host galaxy contribution to the monochromatic luminosity into account, leads to a slightly different expression for the black hole mass, but does not alter our results significantly.

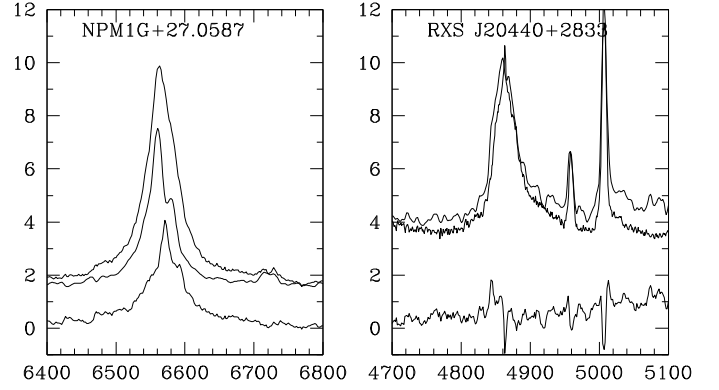


Fig. 4. Profile variations for two objects, NPM1G+27.0587 and RXS J20440+2833, H α and H β regions respectively (see the text). The original spectra taken at two different epochs (Table 1) are shown as thick lines and the difference spectrum, as a thin one. Artificial offset may be added for a better presentation.

The last column of Table 3 shows the type of the object (e.g. S1/2, LINER, HII galaxy), based on the presence of a broad component or different narrow line (diagnostic) ratios. Although all the objects were initially classified as S1 types (according to VCV03), our spectra clearly indicate that some of them should probably be classified as S1.9/S2/LINER types (Table 3). To be able to distinguish between the narrow-line objects (e.g. S2, LINER and HII galaxies), we employed diagnostic diagrams (Veilleux & Osterbrock 1987; Véron-Cetty et al. 2001), where ratios like $[OIII]\lambda 5007/H\beta_{narrow}$, $[SII]\lambda 6716+6731/H\alpha_{narrow}$, $[OI]\lambda 6300/H\alpha_{narrow}$, etc. are used to separate objects of different types. We also checked the place on these diagrams of some broad wing objects (i.e. S1.8, S1.9) to look for possible LINER 1 type objects (i.e. LINERs with broad line wings), whose existence has been suggested by Ho et al. (1997a,b). We found, however, no clear evidence of such objects in our sample. See Sect. 3.5 for classification details.

3.3. Profile variations

A few objects have been observed spectroscopically more than once (Table 1). Comparing spectra from different epochs could be of particular interest to study the line profile variability of the objects. Furthermore, a visually bright, variable quasar can be a good candidate for eventual future reverberation-mapping campaigns. Profile changes are shown in Fig. 4 for two objects (NPM1G+27.0587 and RXS J20440+2833). The H β region was preferred, unless the quality of the spectra was not good enough in this region, so H α was used instead. Before subtracting, the spectra were scaled in such a way that no sign of the narrow lines to be seen in the resulting spectrum ([OIII] for H β and [SII] for H α). This guarantees that variable atmospheric conditions, which may change artificially the continuum, will not alter the broad profiles. One sees that, while for RXS J20440+2833 the results are not conclusive enough in terms of profile variations, NPM1G+27.0587 clearly shows a change in H α flux (by about 50%), making it quite suitable for reverberation mapping studies in the future.

3.4. Reddening

Since most of our spectra cover both H β and H α regions, we explored the connection between the Balmer decrement and the optical continuum slope in more details in order to study how the reddening affects the central continuum source and the BLR,

especially taking into account that many of our objects are heavily reddened. Figure 5 shows the V-band absorption affecting the BLR, assuming a canonical intrinsic broad line decrement of 2.8, A_V^{BLR} , vs. the continuum absorption, A_V^{cont} , assuming an intrinsic optical slope of $a_V^{\text{intr}} = 2$, where $F_\nu \sim \nu^{a_\nu}$. The observed continuum slope, a_V^{obs} , was measured between 5200 and 6200 Å, Table 1, and then the absorption can be roughly estimated as $A_V \approx 0.83(a_V^{\text{intr}} - a_V^{\text{obs}})$. The choice of a_V^{intr} was made based on the steepest objects (presumably unabsorbed) that we have in our sample (see also Dietrich et al. 2005, whose data is also added in Fig. 5), but our main findings are not affected by the exact value of a_V^{intr} (see also Sect. 4.3). As one may also see from Fig. 5, A_V^{cont} and A_V^{BLR} correlate rather well (Pearson correlation coefficient of 0.71), especially taking possible uncertainties in both directions. It is interesting to note, however, that the slope of the best fit is not 1 as one can expect, but close (within errors) to 2 instead; Sect. 4.3 discusses this issue.

3.5. Individual notes

Since most of the objects from our sample have not been previously discussed in the literature, here we consider some of their most interesting features.

2MASX J005050+3536. A spectrum of this object of slightly worse spectral resolution is published by Wei et al. (1999). Both spectra look quite alike in terms of broad line profiles, continuum level, etc. See also Grazian et al. (2000) for an additional spectrum.

MCG+08.17.060. Due to unfavorable atmospheric conditions, the quality of this spectrum is rather low and no attempts to subtract iron or host galaxy have been made. For the same reason, the $H\beta$ profile has not been measured and no M_{BH} estimated. The active galaxy is a component of an interacting pair. Kollatschny et al. (2008) have recently published a better spectrum of this object. Their measurement of $FWHM(H\alpha)$ is similar within the errors. The FIRST radio emission (≈ 18 mJy; Bauer et al. 2000) is not sufficient to place the object into the radio-loud category.

RXS J11401+4115. This object again has a rather noisy spectrum, so the FeII measurements are rather insecure. The SDSS spectrum of this object clearly indicates an S1 galaxy with strong FeII emission, but also potentially strong host galaxy continuum. The Strasbourg CDS Simbad database gives $V = 13.8$ for this object, but our observations suggest the object should be at least a magnitude fainter. Even fainter, the continuum appears to be on the SDSS spectrum. The level of detected radio emission (2 mJy) from this object places it well within the radio-quiet category.

PG 1211+143. This is a well known NLS1 quasar. Reverberation mapping studies (Peterson et al. 2004) infer a central mass, rather similar to ours (Table 1) even though $FWHM(H\beta)$ and the continuum level are slightly different (see also Kaspi et al. 2005).

RXS J12308+0115. This is a radio-quiet NLS1 object with strong FeII. Located very close to 3C 273 in the sky. In a recent paper, Landt et al. (2008) report $H\beta$ flux and continuum measurements twice as high as ours, a possible indication of significant variability, yet the width of the line remains the same within $\sim 4\%$.

RXS J16312+0955. This compact RASS source has a low-dispersion spectrum, published in Grazian et al. (2000).

RXS J17233+3630. This active galaxy is a part of an interacting pair. De Grijs et al. (1992), published a spectrum

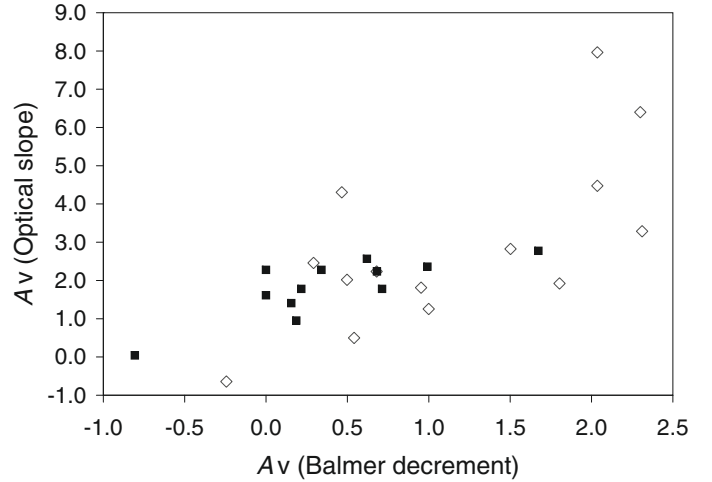


Fig. 5. A comparison of the absorption indices (A_V), calculated through the Balmer decrement of the broad components and the optical slope, assuming an unabsorbed slope of $a \approx 2$ (see the text). The open symbols represent our data, the filled ones are NLS1's from Dietrich et al. (2005). The typical error of A_V should be in order of ± 0.5 in both directions.

of the objects of a slightly worse quality. The detected radio-emission of 3 mJy (Wadadekar, 2004) classifies the object as radio-quiet. The ratio of [SII] $\lambda 6716$ and [SII] $\lambda 6731$, a doublet separated in this spectrum, is a good density indicator and implies $N_e \sim 1000 \text{ cm}^{-3}$ for the NLR (e.g. Osterbrock, 1974). Due to the variable atmospheric conditions when the spectrum was taken, all the continuum and flux measurements (Tables 2 and 3) should only be considered as a rough estimate.

IRAS 18423+2201. Moran et al. (1996) published a spectrum of this object, where the broad component of $H\beta$ appears much stronger and probably broader. The low level of radio emission (2 mJy; Condon et al. 1998) classifies the object as radio quiet. Although the broad lines are clearly seen in the spectrum, the place of this object on the diagnostic diagrams (e.g. [OIII]/ $H\beta$ vs. [OI]/ $H\alpha$ and [OIII]/ $H\beta$ vs. [SII]/ $H\alpha$) appears to be on the border between the Seyferts and the LINERs.

NPM1G+27.0587. This is an IRAS object behind the Milky Way with a published spectrum around $H\alpha$ (Takata et al. 1994).

RXS J20440+2833. This is an object behind the Milky Way (Motch et al. 1998)

NPM1G-05.0589. This spacially extended object (galaxy) does not show broad lines with the exception of a very broad wing of $H\alpha$. The host galaxy determines the optical continuum entirely. All ratios – [OIII]/ $H\beta$, [OI]/ $H\alpha$, [SII]/ $H\alpha$, and [NII]/ $H\alpha$ indicate clearly Seyfert galaxy, so we classified the object as S1.9. The ratio of the [SII] doublet lines implies an NLR density of about 700 cm^{-3} . A low-dispersion spectrum of the object is also available in Grazian et al. (2002).

RXS J21240-0021. It is another object, where the host galaxy continuum appears to significantly dominate the AGN continuum; however, broad wings ($H\alpha$, $H\beta$, HeII $\lambda 4686$) are clearly seen. The place of the object on the diagnostic diagrams is on the border between Seyfert-III-LINER galaxies, suggesting the possible composite nature of its spectrum. Most likely the object should be classified as a low luminosity S1 galaxy. The [SII] ratio of 1.15 implies an NLR density of about 600 cm^{-3} . The SDSS spectrum and a low dispersion one in Grazian et al. (2002) are available as well.

NPM1G+24.0470. This spacially extended object probably also has a composite nature, being on the border between

Seyfert-III-LINER regions on the diagnostic diagrams. Weak broad lines are probably present in the $H\alpha$ profile, which technically means the object should be classified as S1.9. The host galaxy is mostly contributing to the continuum. The [SII] ratio of 0.88 implies a density of about 2000 cm^{-3} .

RXS J21592+0952. This is a broad-line object that appears to have a double-peaked profile (e.g. Eracleous & Halpern 1994; Strateva et al. 2003), clearly seen in $H\beta$. The object was also observed by Wei et al. (1999), whose spectrum reveals stronger broad emission. It is a weak radio source, well within the radio-quiet range.

RXS J22027-1304. This object is a spatially extended galaxy. On the diagnostic diagrams, it stays well inside the Seyfert region and should be classified as either S1.5 or S1.8. No signatures of FeII are seen.

RXS J22160+1107. Here the continuum is largely dominated by the host galaxy. Diagnostic ratios suggest the Seyfert nature of the object; and taking into account the weak broad wings, we classify the object as low-luminosity S1 or S1.8.

RXS J22287+3335. This is a rather bright S1 type object with moderate FeII emission. Wei et al. (1999) published a spectrum of slightly worse quality, on which the object, interestingly, appears to be much fainter (≈ 5 times in terms of the continuum level), indicating possibly strong variability and $H\beta$ appears to be narrower there as well.

NPM1G-04.063. This is a very interesting object with lines of significant width but continuum – clearly dominated by the host galaxy. If the starlight correction is correct, the intrinsic AGN continuum appears to diminish towards the shorter wavelengths, down to zero around $H\beta$, which suggests significant absorption between the continuum source (e.g. the accretion disk) and the BLR. Some of this absorption may also be associated with the BLR itself, due to the quite large Balmer decrement for the broad H components (≈ 6).

4. Discussion

4.1. Line profiles

While fitting the line profiles, we did not follow a multi-Gaussian/Lorentzian approach, because two or three, say Gaussians, fitting the broad profile can hardly have any physical meaning (in the sense that they can hardly represent different physical regions). Furthermore, each complex profile can be fitted to the desired level of accuracy with a large enough number of functions; however, the fit is not necessarily meaningful, not even unique in terms of degeneracy. We have to mention, though, that multi-Gaussian fits are often used by other authors, mainly for the sake of a better description of the profile (Warner et al. 2003; Popovic et al. 2004; see also Bachev et al. 2004, and the discussions there). The impression that the narrower profiles are better described by a Lorentzian, not Gaussian, (Sulentic et al. 2000; Véron-Cetty et al. 2001), however, may have no physical meaning, since it is not clear to what extent *FWHM* measures physically meaningful velocity-related quantity in a complex profile. A more adequate measure would be the second moment of the profile (the velocity dispersion), which is infinity for a Lorentzian. Even if this were not the case (say for a modified Lorentzian), a Lorentzian-type profile typically has higher dispersion than a Gaussian of the same *FWHM*. Therefore, it would be biased to conclude that from two profiles of the same velocity dispersion, the narrower one (in terms of *FWHM*!) resembles a Lorentzian more than a Gaussian.

4.2. The obscured AGN population

Finding more and more rather bright (i.e. 14–15 mag) AGN shows clearly that color-based quasar surveys, like PG for instance, are not complete. This is obviously so, because the quasars are very different, not only in terms of emission-line, radio, X-ray properties, but also in terms of the optical continuum slopes, which ultimately define the colors. The intrinsic quasar continuum obviously may not only depend on the exact accretion solution type and accretion parameters, such as black hole mass, accretion rate, accretion disk inner and outer radii, etc. The continuum shape and level can be severely altered by the host galaxy contribution and the internal absorption, and we found evidence of both in the sample we studied. The reddened continuum slopes can eventually explain why such bright object are not discovered earlier as AGN. On the other hand, a correct assessment of the intrinsic λL_{5100} is tightly linked to the estimates of the bolometric luminosity and the central mass through empirical relations (the often used $L_{\text{bol}} \approx 9\lambda L_{5100}$, see also Sect. 3.2). It also affects other characteristics, like optical-to-X-ray slopes, radio-loudness (e.g. Ho 2002), line equivalent widths, etc. Therefore, one needs to know the intrinsic quasar continuum for accurately computing the accretion parameters and exploring correlations like EV1. For instance, large quasar samples with well-studied spectral properties (e.g. Boroson & Green 1992; Marziani et al. 2003) rarely make an attempt to assess the real continuum level through correcting for the starlight and the intrinsic absorption, presuming that their contribution is small. It may indeed be so for some bright quasars, but as we see here (see also Dietrich et al. 2005), this is not always the case. Furthermore, dangerous biases may be introduced – the host galaxy will contribute to the continuum mostly for a lower-luminosity AGN, and the internal absorption may be connected with the orientation if associated with a thick torus, and the broad-line widths if the torus is coplanar to a flattened BLR. It should be easy to correct for the host galaxy in a spectrum of relatively good quality. Unfortunately, the host galaxy contribution depends largely on the slit size, seeing, etc. and therefore may differ significantly from observation to observation and from instrument to instrument. This also can have consequences for the reverberation-mapping studies, in the sense that a variable starlight contribution can at least add some extra noise in the cross-correlation function between the continuum and the lines.

4.3. Where is dust located?

In this paper we did not make any attempt to correct the continuum for the intrinsic absorption; nevertheless, the latter is rather apparent for objects like NPM1G-04.063, where quasar continuum around $H\beta$ appears to be totally absorbed.

One may make an attempt to take this absorption into account either by measuring the continuum slope or by comparing the Balmer decrement with the theoretical expectations (i.e. case “B” recombination). Both approaches may lead to incorrect results, however. The slope can, of course, vary for a number of reasons and the dust reddening is only one of them. On the other hand, the Balmer decrement can also be misleading since we do not know exactly where the absorbing material is. If it is uniformly spread outside the NLR, then such a correction can in principle be performed, and the absorption will presumably equally affect the NLR, the BLR, and the central continuum source. Different Balmer decrements for the broad and the narrow lines, as seen in some of the objects in our sample, suggest this may not always be the case. If the absorbing material

is at least outside the BLR, then the broad line decrement could be used to assess the unabsorbed continuum level and shape. Unfortunately, we cannot rule out the presence of absorption inside the BLR inner radius (or mixed with the BLR matter), in which case the continuum will be affected, but indications for this might not be seen from the emission line ratios.

An important clue to solving the problem of the dust location can be provided by the relation between the optical continuum slope and the Balmer decrement. We found a correlation between the slope and the decrement absorption indices (Fig. 5). One is to note, however, that the wide range of both – Balmer decrements and continuum slopes observed in our rather small sample is not typical. For instance, Dong et al. (2008) report a much narrower range of Balmer decrements for a large sample of SDSS quasars – typically between -0.5 and 1 , and slopes – between 1.2 and 2.2 (both expressed in terms of A_V). Their sample is, however, blue-selected and is obviously biased against the obscured AGN population. This may explain why Dong et al. (2008) find no correlation between the decrements and the slopes in their sample.

If further confirmed, the relation $A_V^{\text{cont}} \simeq 2A_V^{\text{BLR}}$ (Sect. 3.4) suggests that the reddening on average affects the center more than the BLR. Even adopting a different value for the intrinsic continuum slope (like the often assumed $a_V^{\text{intr}} \simeq 1$) only offsets the relation between the Balmer decrement and the continuum slope, but does not change the slope of the relation. Adopting a different value for the intrinsic Balmer decrement (e.g. 2.8 instead of 3.1) has a similar effect. One way to explain this result is to assume that the dust is associated with the BLR matter and the reddening affects the continuum coming from inside more than the BLR emission itself. Another explanation would be that the dust is much farther outside the BLR (probably associated with a torus), and we see the active nucleus just above the torus edge. As a result, the central continuum can appear heavily absorbed, but the BLR light is not that affected, taking their different spatial dimensions into account. If the latter explanation turns out to be the correct one, studying the line profiles of such partially obscured objects can be important for probing the exact BLR geometry and kinematics, since knowing the torus opening angle, the inclination can be estimated for these objects.

5. Summary

We obtained optical spectra of bright AGN, mostly recently discovered as X-ray sources. Different emission line parameters were measured, including important EV1 characteristics, such as $FWHM(H\beta)$ and R_{Fe} . Black hole masses were estimated through an empirical relation, invoking the line width and the continuum level. We show that the host galaxy contribution to the continuum may be non-negligible in many cases, often about 50% or more. The host galaxy, as well as the intrinsic absorption, can significantly alter the continuum slope, offering a possible explanation why such bright object have not been discovered earlier as AGN. Based on the line widths and emission line ratios, we provided a Seyfert type classification of each object, showing that, although previously listed as S1 types, some objects should

be considered at least S1.9, if not S2 or LINERs. Finally, taking into account that many of our objects are heavily reddened, we discuss different possibilities about the exact location of the absorbing material.

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