

Eruptive activity of Z And in 2000 – 2010*

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Abstract. Results of the study of the prototype of the classical symbiotic stars Z And during its last active phase 2000 – 2010 when it experienced a series of six optical outbursts are presented. High-resolution spectra obtained during the first and fourth outburst, which was the strongest one, have been analyzed. These data are compared with results of theoretical computations. The comparison provides information about the behaviour of the system during the entire active phase rather than during an individual outburst. In particular it was found fundamental difference between the first outburst which opened the active phase and the recurrent outbursts - namely, the presence of bipolar collimated optical outflow during some of the recurrent outbursts. A scenario that can explain all the spectroscopic phenomena observed during this active phase as well as previous active phases of Z And is proposed. The possibility to use this scenario for explanation of the line spectrum of other classical symbiotic stars during their active phases is motivated.

Key words: binaries: symbiotic - stars: activity - stars: winds, outflows - stars: individual: Z And

Еруптивна активност на Z And през 2000 – 2010 г.

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Представени са резултати от изследването на прототипа на класическите симбиотични звезди Z And по време на последната ѝ активна фаза в периода 2000 – 2010 г., когато претърпя серия от шест оптически избухвания. Анализирани са спектри с високо разрешение получени по време на първото и четвъртото от избухванията, което беше най-силно. Тези данни са сравнени с резултати от теоретично моделиране. Сравнението дава информация за поведението на системата не само по време на отделните избухвания, но и през активната фаза като цяло. В частност, намерена е фундаментална разлика между първото избухване, инициращо активната фаза и повторните избухвания, а именно – присъствието на оптическо биполарно колимирано изтичане на маса по време на някои от повторните избухвания. Предложен е сценарий, в рамките на който могат да се обяснят всички спектрални характеристики, наблюдавани както през тази активна фаза, така и по време на предишните. Мотивирана е възможността този сценарий да се използва за обяснение на линейния спектър и на други класически симбиотични звезди по време на активни фази.

Introduction

Symbiotic stars are interpreted as interacting binaries consisting of a cool visual primary and a hot compact secondary component accreting matter from the atmosphere of its companion. Their spectral variability is determined from the orbital motion and the outburst events of the hot component.

* Based on observations collected at the Rozhen National Astronomical Observatory, Bulgaria

These events are often accompanied by intensive loss of mass in the form of optically thick shells, stellar wind outflow and bipolar collimated jets. The stellar wind is possible to be bipolar and highly collimated too. The ejected material forms complicated flow structure in the space between the components, whose elements determine the observed properties of the star. The interacting binary Z And is considered as a prototype of the classical symbiotic stars. It consists of a normal cool giant of spectral type M4.5 (Murset & Schmid 1999), a hot compact component with a temperature of about $1.5 \cdot 10^5$ K (Sokoloski et al. 2006) and an extended circumbinary nebula formed by the winds of the stellar components and partly photoionized by the compact object.

Z And has undergone several active phases consisting of repeated optical brightenings with amplitudes up to 2-3 mag and characterized by intensive loss of mass (Swings & Struve 1941; Boyarchuk 1967; Fernandez-Castro et al. 1995; Bisikalo et al. 2006; Skopal et al. 2006; Sokoloski et al. 2006; Tomov et al. 2008; Skopal et al. 2009; Tomov et al. 2010). The last active phase of Z And began at the end of August 2000 (Skopal et al. 2000) and continued up to that time including six optical brightenings. The maxima of the light during the active phase were in December 2000 ($V \sim 8^m8$), November 2002 ($V \sim 9^m8$), September 2004 ($V \sim 9^m1$), July 2006 ($V \sim 8^m6$), January 2008 ($V \sim 9^m5$) and January 2010 ($V \sim 8^m4$) (Skopal et al. 2009; Skopal, private communication; data of AAVSO). High resolution optical data indicating highly different physical conditions (velocity, density, temperature) in the line emission regions were obtained during this active phase (Skopal et al. 2006; Sokoloski et al. 2006; Tomov et al. 2007; Tomov et al. 2008; Skopal et al. 2009; Tomov et al. 2010).

The current study is based on observations acquired during the first and fourth outbursts. The main features of the flow structure in the system during the active phase indicated by our data can be summarized as follows:

1. We suppose that an accretion disk is present in the system. Only in this way we can explain the presence of the He II λ 4686 line in the spectrum and its behaviour during the first outburst (Tomov et al. 2008).
2. Stellar wind was observed in the system during all the outbursts. Two wind components with different velocity regimes were observed simultaneously – wind propagating at a moderate velocity, $0 - 100 \text{ km s}^{-1}$ (Sokoloski et al. 2006; Skopal et al. 2006, Tomov et al. 2008), and wind with a high velocity, in excess of 500 km s^{-1} (Skopal et al. 2006, Tomov et al. 2008).
3. Bipolar collimated outflow appeared during the fourth outburst. In the beginning of July 2006 the H α line had an absorption component shifted by -1400 km s^{-1} from its center. It went into emission and later, in July – December 2006, the spectrum contained additional emission components on either side of the central peak corresponding to velocities of $1200 - 1500 \text{ km s}^{-1}$ (Burmeister & Leedjarv 2007; Tomov et al. 2007; Skopal et al. 2009). The H β line had the same features (Tomov et al. 2007; Skopal et al. 2009). Both the emission and absorption high-velocity components are assumed to be formed in a bipolar collimated outflow (Tomov et al. 2007).

The main aim of the current study is to suggest a model for the flow structure in the system to explain all of these features. We believe that it will be helpful in explaining not only the data acquired during this active phase, but also all spectroscopic data taken during earlier phases, those following 1939, 1960 and 1984 and possibly the activity of other classical symbiotic stars.

1 The model of the flow structure in Z And

According to the theoretical models (Bisikalo et al. 2002; Mitsumoto et al. 2005; Bisikalo et al. 2006) an accretion disk is formed around the compact object in the Z And system when the wind of its giant has quiescent parameters. The radius of the disk is about $50 \mathcal{R}_{\odot}$ for a wind velocity of 20–25 km s^{-1} and thus the outer part of the disk is optically thin. The mass of the disk is estimated as product of one quarter of the mass-loss rate of the giant $\sim 2 \times 10^{-7} \mathcal{M}_{\odot} \text{yr}^{-1}$ (Fernandez-Castro et al. 1988), and the typical time interval between the active phases – about 10 years. Based on the size and the mass of the disk of $\sim 5 \times 10^{-7} \mathcal{M}_{\odot}$ we can assume that its inner region is optically thick in the quiescent state of the system.

To change the system from quiescent to active state a sufficiently large increase of the accretion rate is needed. Accretion of a considerable fraction of the disk’s mass requires an outburst to develop even in the combined model where the increased nuclear burning rate is taking into account (Bisikalo et al. 2006). Maximal increase of the accretion rate is possible in the framework of the mechanism proposed by Bisikalo et al. (2002) and Mitsumoto et al. (2005). According to this mechanism, even a small increase of the velocity of the wind of the donor is sufficient to change the accretion regime. During the transition from disk accretion to accretion from the flow, the disk is partially disrupted and the increased velocity of the wind causes falling of the material of the disk onto the accretor’s surface. However, even in this case a considerable amount of mass (up to 50 – 80 per cent) stays in the disk. Then massive accretion disk exists in the system during the active phase too.

During the outburst the high velocity wind of the hot compact component collides with the accretion disk, its velocity decreases near the orbital plane but does not change at higher stellar latitude. A consequence of this is that two different kinds of stellar wind are observed – P Cyg wind with a low velocity and an optically thin high velocity wind. As a result of the collision the density of the wind close to the orbital plane increases and the level of the observed photosphere locates at a larger distance from the star. Thus, an optically thick disk-like shell forms, which occults the compact object. Since its effective temperature is lower it is responsible for the continuum energy redistribution and the increase of the optical flux of the star. The interaction of the wind with the disk is equivalent to interaction of two winds and thus can form shocked region of collisional ionization with temperature of up to 10^6 K (Fig. 1) (Nussbaumer & Walder 1993; Bisikalo et al. 2006).

During the active phase the wind from the hot component “strips” the accretion disk carrying some part of its material in the circumbinary envelope. After each outburst some part of this material does not leave the potential

well and after the cessation of the wind begins to accrete again. Because of the initial amount of angular momentum a disk-like envelope extending to larger distance from the orbital plane than the accretion disk itself forms (Fig. 2). The existence of a centrifugal barrier creates hollow cones around the axis of rotation with small opening angle (Icke 1981; Blandford & Begelman 2004).

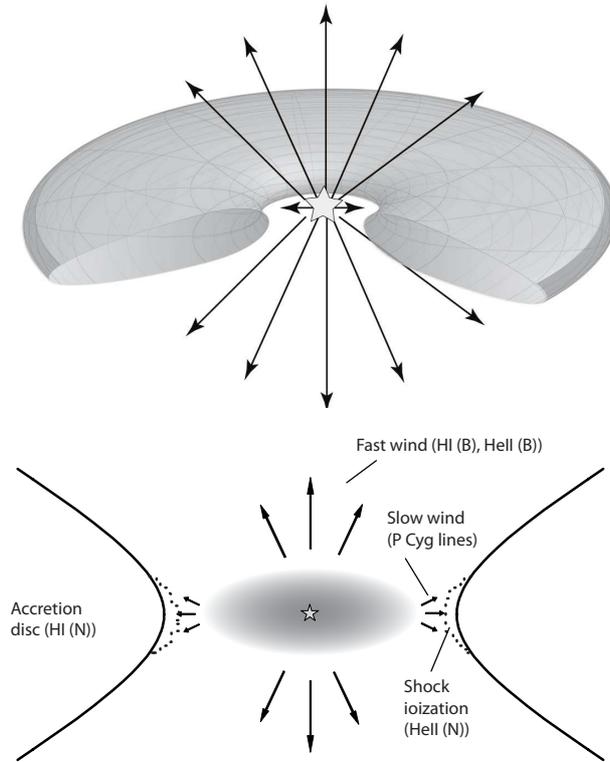


Fig. 1. Above panel: Schematic model of the region around the hot component during the first outburst. Below panel: The same, but in the plane perpendicular to the orbital plane where the emission regions are shown. (From spectroscopy presented in Tomov et al. (2008).)

The disk-like envelope does not exist during the first outburst of the active phase. During the recurrent outbursts the extended disk-like envelope can collimate the wind which escapes the compact object predominantly via the two hollow cones (Fig. 2). The bipolar outflow can be observed as high-velocity satellite components situated on either side of the mean peak of the

line. Their presence in the spectrum depends on the density of the disk-like envelope (i.e. on the system's activity during its previous phases) and on the wind intensity. They can appear only if the density of the disk-like envelope is sufficient to provide collimation and the mass-loss rate of the compact component is high enough. According to our model these spectral features can be observed during outbursts accompanied by loss of mass at high rate and preceded by similar strong outbursts.

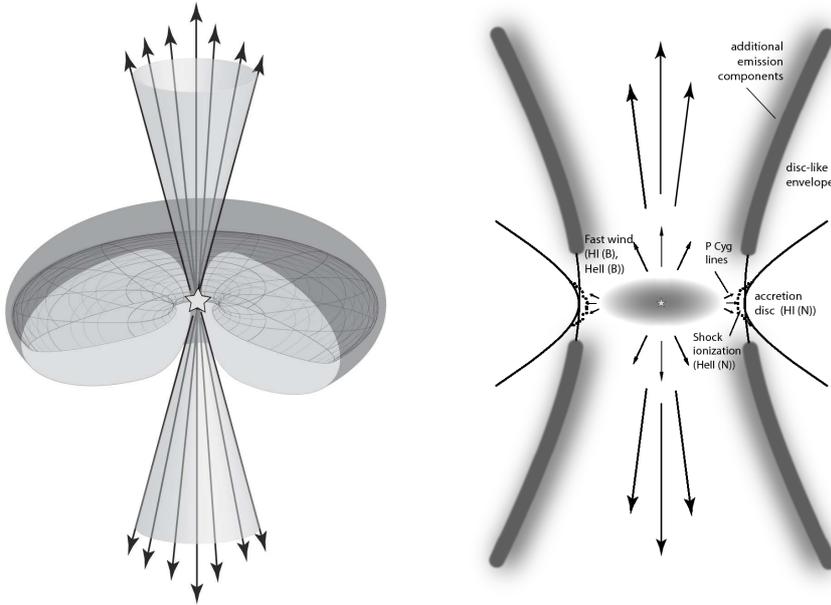


Fig. 2. The same like in Fig. 1 but during recurrent strong outburst.

According to the current theory the existence of bipolar collimated jets is supposed to be due to presence in the system of magnetized disk which transforms the potential energy of the accreting material into kinetic energy of the outflowing jets. This means that accretion must be realized to provide existence of the jets. In the spectrum of Z And during its 2006 outburst, however, different indications of stellar wind presented along with satellite emission components. If we believe that “traditional” jets were observed, it is necessary to explain accretion onto the star at the time when a strong mass outflow from its surface was observed. On the other hand, the proposed model of collimated stellar wind makes it possible to explain all features of the observed spectrum without any important contradictions.

An other question related to the nature of the collimated outflow is that about the orbit inclination. The velocity of the radiatively accelerated wind

cannot exceed 3000 km s^{-1} . The highest observed outflow velocity during the 2006 outburst was 1500 km s^{-1} . A preliminary analysis of our spectra obtained during late December 2009 showed velocities of $1700 - 1800 \text{ km s}^{-1}$. Taking into account these velocities we obtain an upper limit of the orbit inclination of about 55° . That is why the view about the nature of the outflow can shed some light on the question about the orbit inclination.

We suppose that the model suggested by us can be used to explain the behaviour of the line spectrum of other classical symbiotic stars during their active phases too since the classical symbiotic stars have the next general characteristics:

1. The mass transfer in the majority of the classical symbiotic stars is realized by means of the stellar wind of the giant. The theoretical computations of Bisikalo et al. (2002) and Mitsumoto et al. (2005) show that for systems with parameters close to those of Z And an accretion disk from wind accretion forms around the compact object. The accretion disk prevents the outflowing material during the active phase and outflow with two-component velocity regime forms.
2. The systems with parameters close to Z And have similar accretion rate of their compact object. The accretion rate determines the regime of hydrogen burning. The view that hydrogen burns in a steady state in the classical symbiotic stars is commonly accepted. When the accretion rate goes above the upper limit of the steady burning range the white dwarf expands which is observed as optical outburst with typical duration of one year. The first outburst opening the active phase can be followed by repeated outbursts. During the repeated outbursts an extended envelope surrounding the accretion disk exists in the system and it can be responsible for the collimation of the stellar wind.

2 Observational evidences in support of the model

2.1 The first outburst

According to our model a geometrically thin accretion disk exists in the system in the beginning of the first outburst and the inner region of this disk can be optically thick.

The presence of a moderate-velocity wind is indicated by the P Cyg profiles of the triplet He I $\lambda 4471$ and He I $\lambda 4713$ lines during the light maximum. The two-component structure of their absorption showed mass outflow from the compact component with a mean velocity of 60 km s^{-1} (Fig. 3) (Tomov et al. 2008).

During the time of the increased light the H γ line had two emission components – a central narrow one and a broad, low-intensity component (Tomov et al. 2008). In November 2000 an absorption feature was probably present to the short-wavelength side of the narrow component, forming a dip in its emission profile (Fig. 4 right panel). The position of the dip corresponded to the red absorption component of the P Cyg lines of He I (Fig. 5). The low velocity of the P Cyg absorption and the dip in the H γ emission profile can

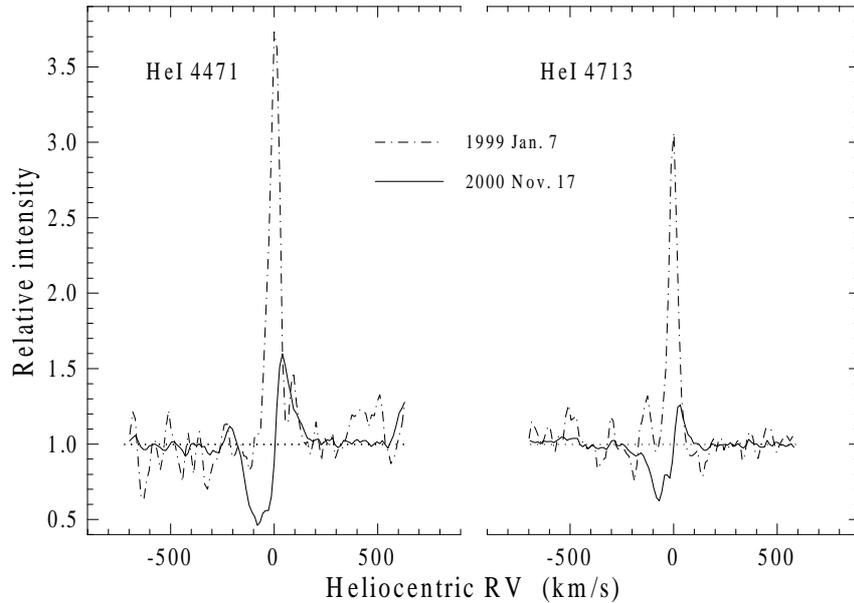


Fig. 3. The profiles of the He I triplet lines in quiescence (January 1999; dot–dashed line) and near the maximum of the first outburst (November 2000; solid line). (From spectroscopy presented in Tomov et al. (2008).)

be explained with supposition for decrease of the velocity of the wind of the compact component after its collision with the disk (Fig. 1b).

The blue wing of the broad $H\gamma$ component was not seen in November 2000 (Fig. 4, right panel). Assuming that the broad component (like the broad component of the He II $\lambda 4686$ line) is emitted by an optically thin high-velocity wind from the compact object, this can be explained with absorption in the stellar wind which is responsible for the P Cyg profiles.

In the quiescent state an additional emission component was present on the short-wavelengths side of the $H\gamma$ line forming a shoulder in its profile. Thus the full width at the zero intensity (FWZI) of the line reached to about 500 km s^{-1} (Fig. 4, left panel). The presence of such a component can be explained with emission by an accretion disc.

The He II $\lambda 4686$ line of Z And was two-component one consisting of a central narrow component and a broad component with a low intensity (Tomov et al. 2008). The broad component was very weak in quiescence and could not be measured accurately. It became stronger during the first outburst and towards the light maximum its FWZI exceeded 1000 km s^{-1} (Fig. 6). We believe that during this period the broad component of the He II $\lambda 4686$ and $H\gamma$ lines emitted in a region of stellar wind with a velocity of 500 km s^{-1} traveling at higher stellar latitudes.

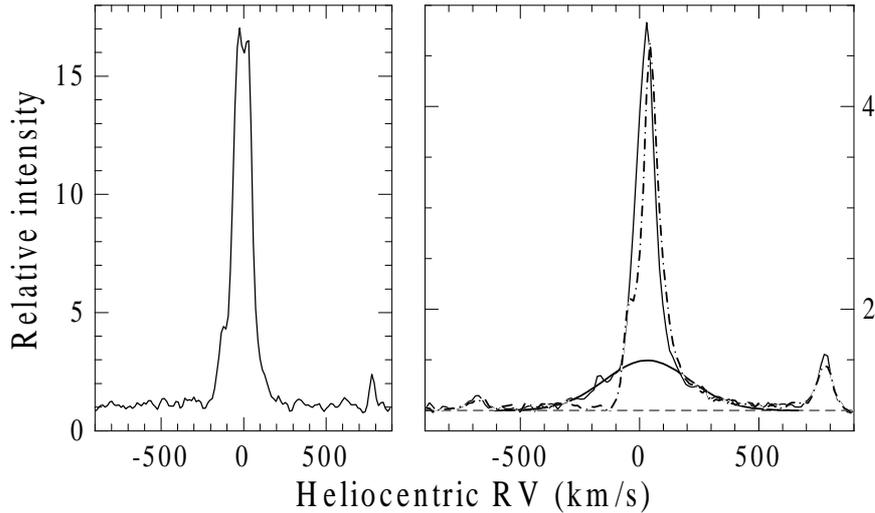


Fig. 4. The profile of the $H\gamma$ line in quiescence of Z And, on January 7, 1999 (left) and at various stages of the 2000 outburst (right). The dot-dashed line in the right panel corresponds to the spectrum on November 17, 2000, and the solid line – to the spectrum on December 5, 2000. The Gaussian fit of the broad component is marked with a thick line. The level of the local continuum is marked with a dashed line. (From spectroscopy presented in Tomov et al. (2008).)

2.2 The fourth outburst

The fourth outburst was the strongest one. The spectral data acquired during this outburst contain indications of all kinds of stellar wind, observed during the first one, and of bipolar collimated outflow in addition. We suppose that a disc-like envelope surrounding the accretion disc and collimating the outflowing material existed in the system at that time and the mass-loss rate was high enough to give rise to satellite line components.

The analysis of our observations of October 2006 with high resolution (0.2 \AA per pixel, as for the first and second outbursts) in the $H\gamma$ and $\text{He II } \lambda 4686$ region of the spectrum of Z And reveals the presence of emission up to 500 km s^{-1} from the centers of these lines (Fig. 7). Skopal et al. (2009) obtained the effective photospheric radius of the outbursting component of $(12 \pm 4) \mathcal{R}_{\odot}$ at the time of the light maximum for a distance of 1.5 kpc. The minimal radius for the distance of 1.12 kpc used by us (Tomov et al. 2008, 2010) is $6 \mathcal{R}_{\odot}$. If the broad emission originated in a Keplerian accretion disk at a distance equal to the radius of the outbursting component, its velocity should be not greater than 140 km s^{-1} , in disagreement with the observations. Like in the case of the first outburst it is natural to suppose that the observed

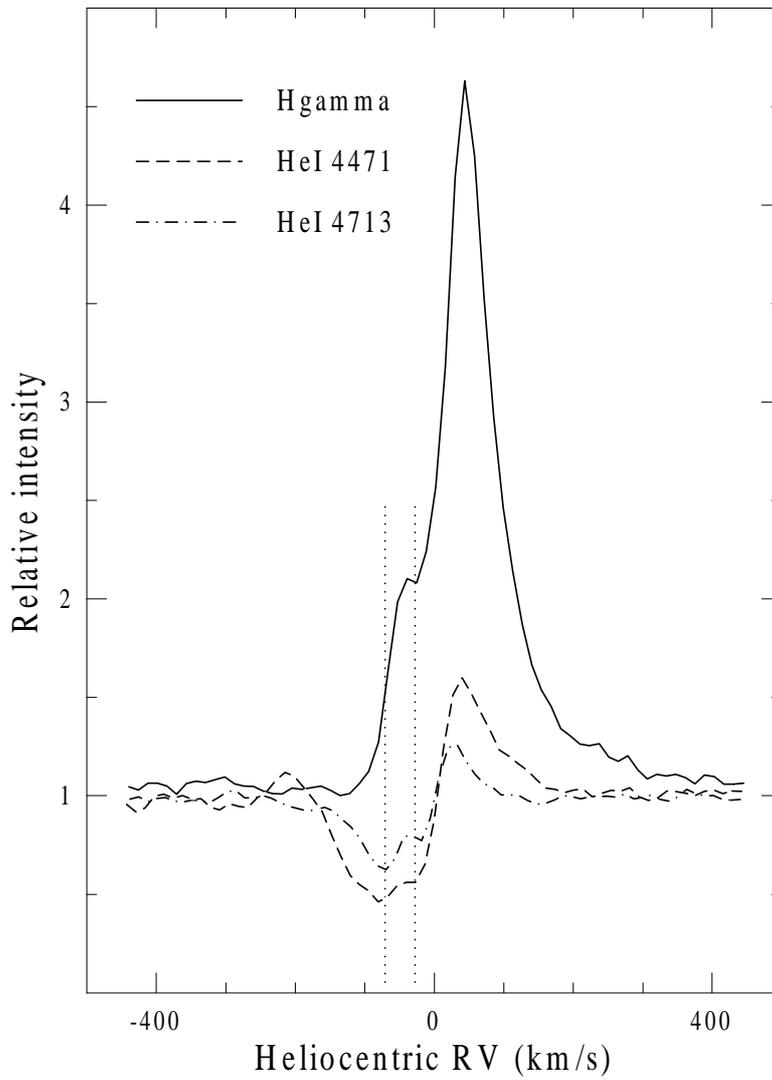


Fig. 5. The profiles of the $H\gamma$, $\text{He I } \lambda 4471$, and $\text{He I } \lambda 4713$ lines obtained on November 17, 2000. The positions of the two components of the He I absorption are marked with vertical dotted lines. (From spectroscopy presented in Tomov et al. (2008).)

broad emission is formed in the high-velocity wind at higher stellar latitudes (Fig. 2). A dip is observed in the blue part of the broad $H\gamma$ emission at a velocity of about 100 km s^{-1} .

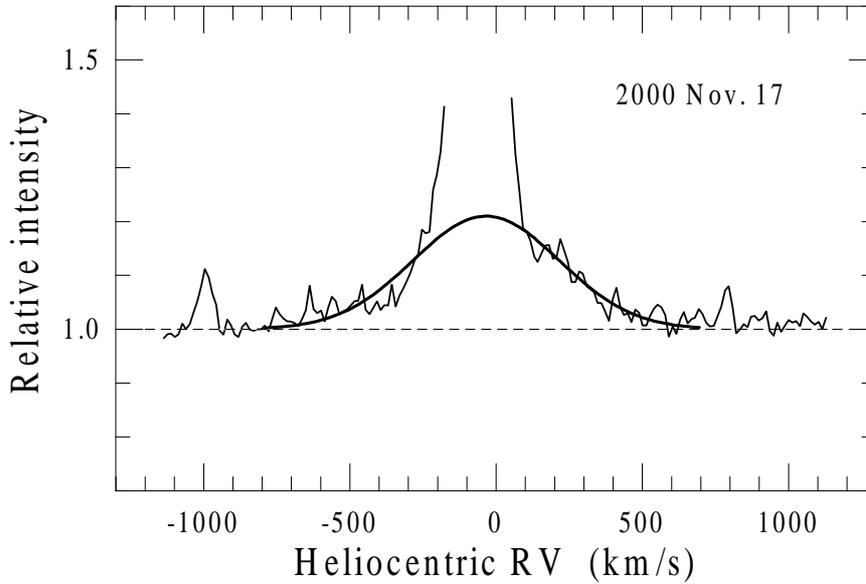


Fig. 6. The wings of the He II $\lambda 4686$ line during the 2000 outburst (November 17). The broad component was visible in this period. (From spectroscopy presented in Tomov et al. (2008).)

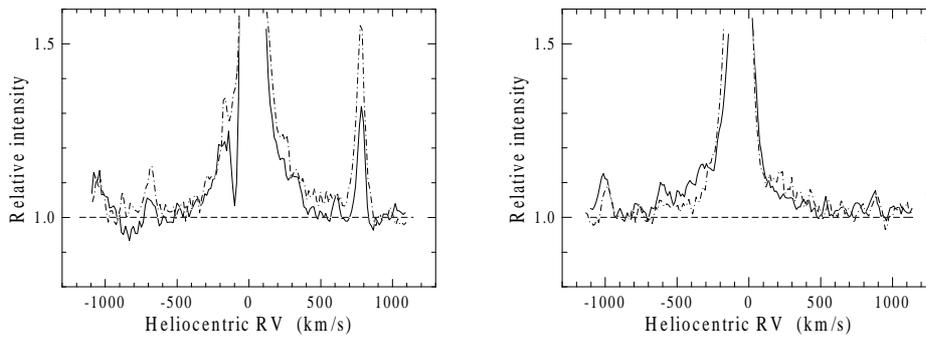


Fig. 7. The broad component of the H γ (left panel) and He II $\lambda 4686$ (right panel) lines during the outbursts in 2000 (December 5; the dot-dashed line) and 2006 (October 4; the solid line). The level of the local continuum is marked with a dashed line.

In our spectra (Tomov et al. 2007), the H α line possessed a strong central component, a broad component (wings) and additional absorption and emission features to either side of the central component (Fig. 8). The central component was single peaked having a shoulder on its short-wavelength

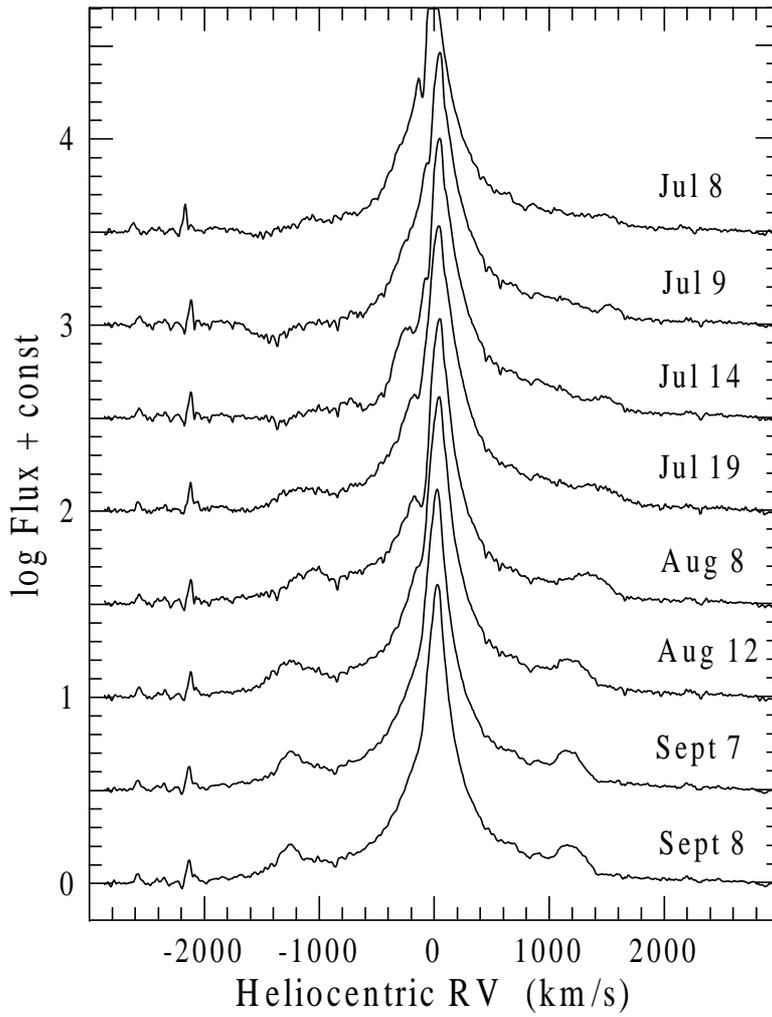


Fig. 8. Evolution of the H α line during the 2006 outburst. (From spectroscopy presented in Tomov et al. (2007).)

side, which was not seen only in the spectra taken in September. A small peak shortward from the central component was seen in the spectrum obtained on August 8, 2006, and the dip before the peak had a velocity about 100 km s^{-1} , in agreement with the results of the analysis of the H γ and He I profiles. Satellite emission components on either side of the central peak

appeared after mid-July (Fig. 8). They are clear signs of the presence of a bipolar collimated wind. According to our model the emitting region is near the cone surfaces (see Fig. 2), where the density of the emerging collimated outflow is the highest.

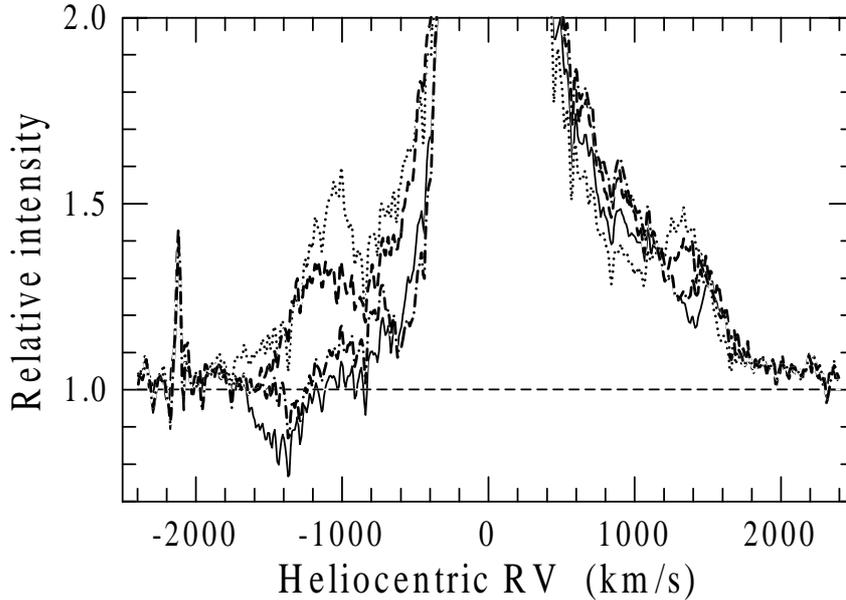


Fig. 9. The variation of the $H\alpha$ line: transition from absorption feature to emission one. The solid line corresponds to the spectrum taken on July 9, 2006; the dot-dashed line to July 14, 2006; the dotted line to August 8, 2006; and the dashed line to September 7, 2006. The level of the local continuum is marked with a dashed line. (From spectroscopy presented in Tomov et al. (2007).)

The first spectra during this outburst taken in July 2006 (Figs. 8 and 9) show a sharp absorption line at a velocity of 1400 km s^{-1} on the short-wavelength side of the central component and only a weak irregularly shaped emission component corresponding to a velocity of about 1500 km s^{-1} on its long-wavelength side. The blue-shifted absorption probably indicates the flow which is projected onto the stellar disk. The evolution of the spectrum in Figs. 8 and 9 shows that the blue-shifted absorption component disappears and an emission appears. Two emission components on either side of the central peak were formed in July. The disappearance of the blue-shifted absorption component and the appearance of emission are probably due to a decrease of the mass-loss rate and/or an increase of the number of emitting atoms in the region of the wind that does not project onto the stellar disk.

During all the outbursts considered by us the $H\alpha$ line possessed broad wings extending to velocities of at least $\pm 2000 \text{ km s}^{-1}$ from its center (Tomov

et al. 2007, 2008, 2010). Skopal et al. (2006) concluded that the H α wings during the outburst state were formed in a high-velocity stellar wind from the compact component. However, as it was noted in Tomov et al. (2010) it is not ruled out that the high-velocity broadening of the H α line was due to radiative damping and the stellar wind contributed to the emission at lower velocities.

3 Conclusions

The comparison of the results of our observations with the results of the gas-dynamical modeling provides possibility to propose a model to explain the behaviour of the Z And system during its whole 2000 – 2010 active phase as well as the differences in its observed properties during different outbursts. The basic points of this model are the following ones:

- ★ An accretion disk exists in the system.
- ★ During the outburst the gas leaving the surface of the compact component at high velocity collides with the accretion disk. As a result its velocity decreases near the orbital plane, but remains unchanged at higher latitudes. Hence, two wind components with different velocities are observed.
- ★ During the outburst the wind from the compact component partially disrupts the disk. Some portion of its mass does not leave the potential well, after the wind ceases it begins to accrete again. Because of the initial amount of angular momentum a disk-like envelope extending to larger distance from the orbital plane forms.
- ★ The existence of a centrifugal barrier creates hollow cones around the axis of rotation which lead to appearance of collimated outflow during the recurrent outbursts. This outflow can be observed as high-velocity components on either side of the main peak of the line.
- ★ Collimated outflow will appear when the density of the disk-like envelope is high enough to provide collimation and the mass-loss rate of the compact component is also high. High-velocity line components thus can be observed during outbursts when the loss of mass is at high rate, preceded by a similarly strong outburst.

The model suggested provides possibility to explain all the spectral features of Z And observed during active stage. The general character of this model permits us to suppose that similar scenario is possible for other classical symbiotic stars too.

Acknowledgments

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Fig. 10. Nikolai Tonov gives a talk